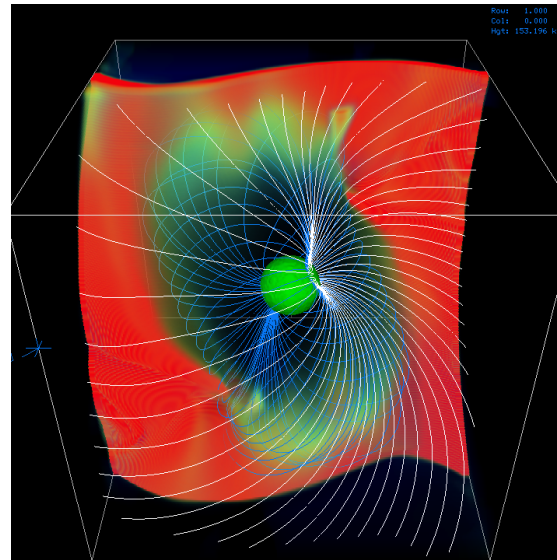
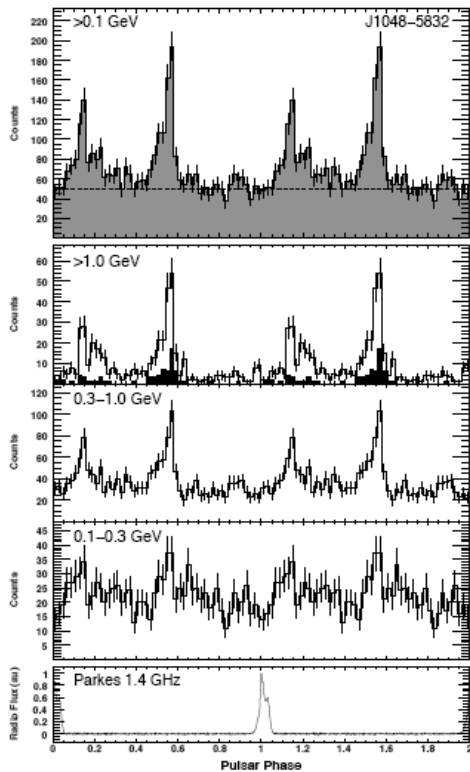
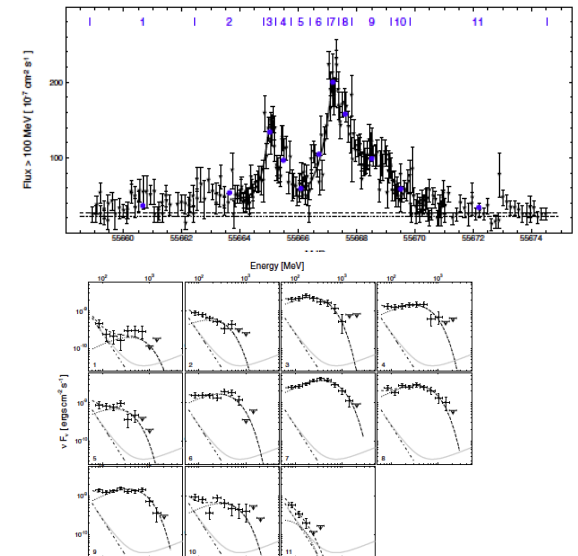


# UHECR from Pulsars/Magnetars based on An "Auroral" Accelerator Model for Gamma Ray Pulsars

Jonathan Arons  
University of California, Berkeley



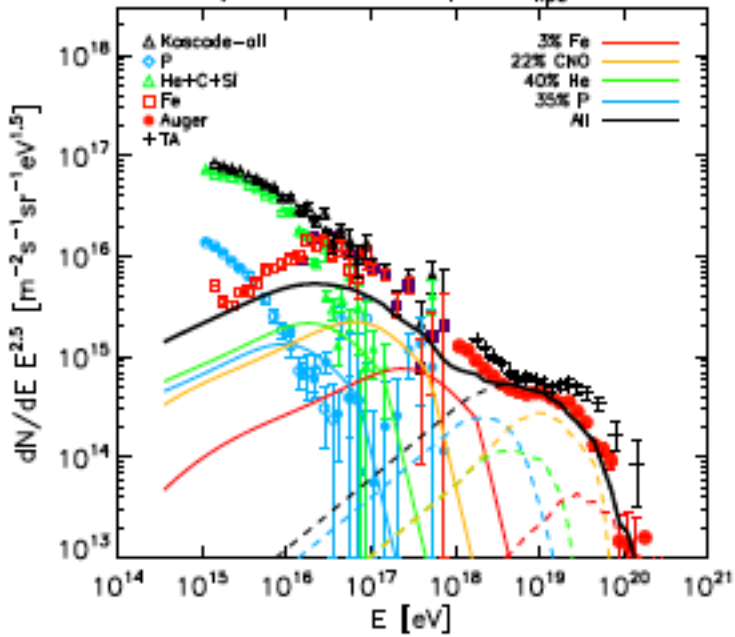
Bai & Spitkovsky 2010



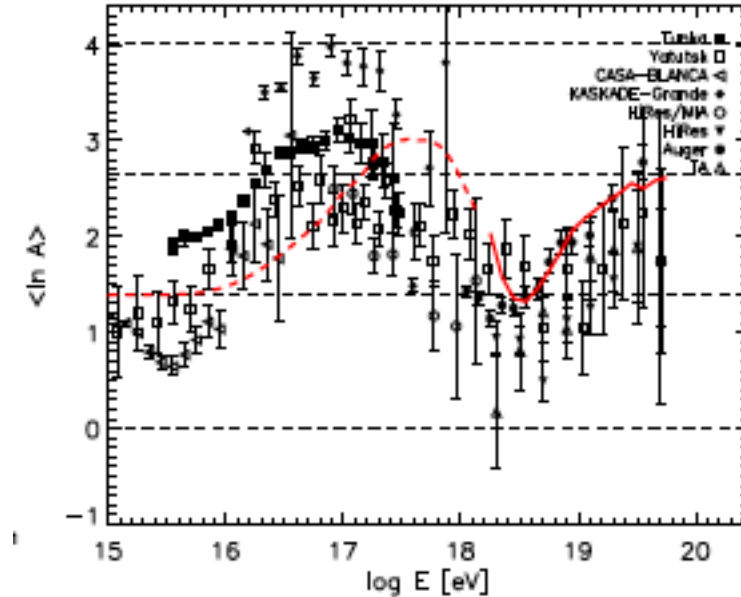
Collaborators: N. Bucciantini, A. Spitkovsky

# Heavy UHECR: a neutron star source?

$\eta=0.3, l_c=20 \text{ pc}, H_{\text{kpc}}=2$



Energy Spectrum  
GZK cutoff or source  
Ends above  $10^{19.5}$  eV



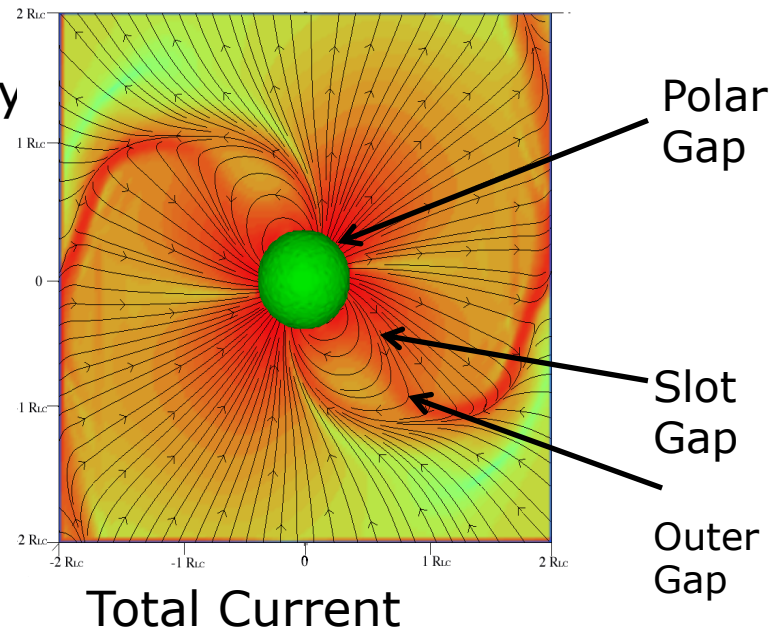
Composition; or hadronic interactions  
in air shower model not understood

Pulsars/Nebulae can accelerate ( $e^\pm$ ) to PeV; have (Fe) crust/ocean

# Magnetospheric Current System Requires Ion Extraction from Star (Atmosphere? Ocean? Crust?)

Aligned/Oblique Rotators structurally similar,  $J_{\text{cond}} + J_{\text{disp}} (=0 \text{ in aligned})$

Spitkovsky's (2006) oblique force free rotator



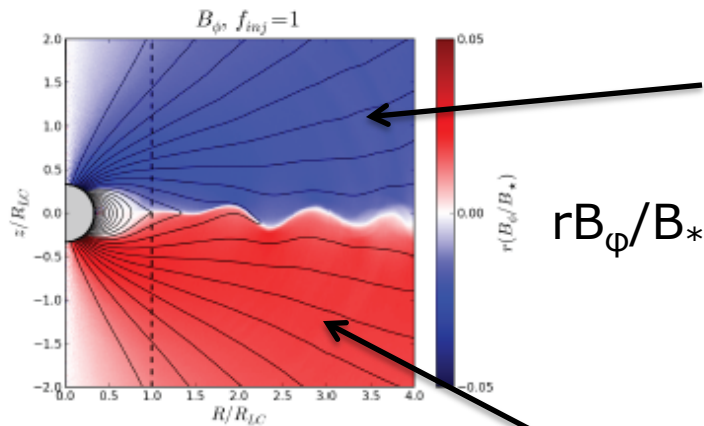
Field Lines (with real open flux)

Gaps = local quasi-**vacuum**  $E_{\parallel}$  **zones inserted by hand into vacuum B** to model gamma ray emission and pair creation – by construction, gaps carry small fraction of total current  $I \Rightarrow L_{\text{gap}}$  small; Accelerate test particles along B rotation  $\rightarrow$  lighthouse  $\Rightarrow$  beamed photons (lighthouse)

$$\dot{E}_R = -I\Omega\dot{\Omega} = k \frac{\mu^2 \Omega^4}{c^3} (1 + \sin^2 i), \quad k = 1 \pm 0.1 \quad = I\Phi, \quad \Phi = \Omega^2 \mu / c^2$$

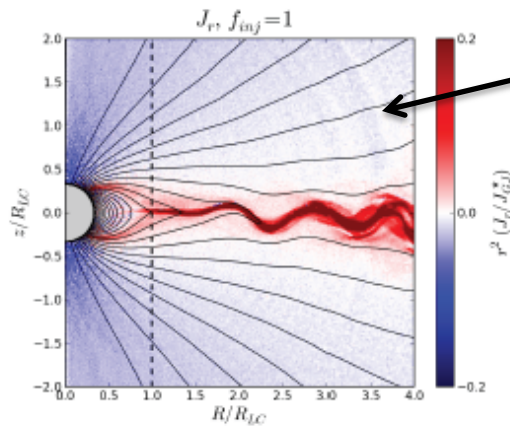
$$i = \angle(\mu, \Omega)$$

**Force Free model has no accelerator: pure MHD** (Alfven's ghost angry): Gap Models with vacuum  $E_{\parallel}$  have too little energy<sub>3</sub>



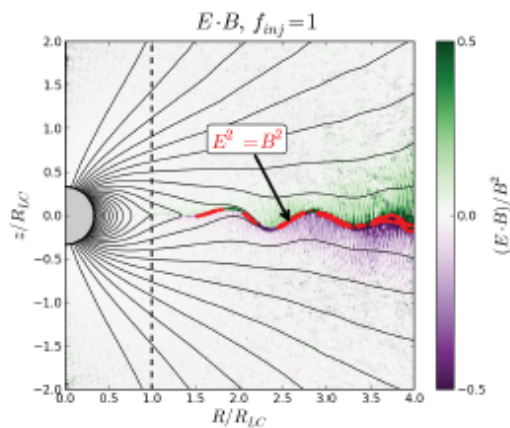
Wind

PIC aligned rotator simulation  
 $e^\pm$  pairs, low altitude pair creation  
 (Cerutti+2015) -  $\sim$ force-free +  
 active current sheet  
 (dynamics, particle acceleration)



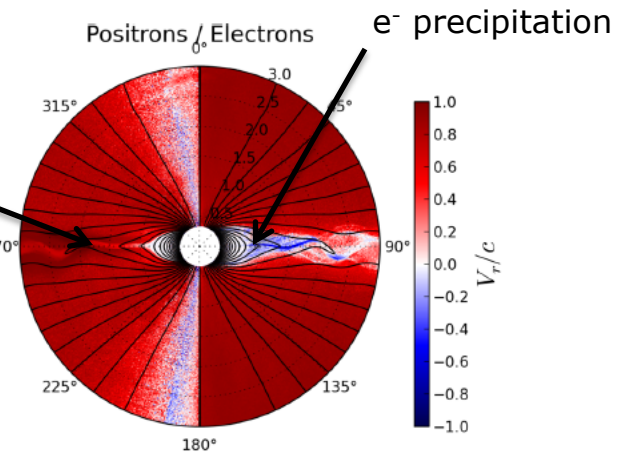
Poloidal field lines

$r^2J_r/J_{GJ*}$  - radial positive current = positive charge  
 outflow (return current in equatorial  
 current sheet) - radial volume current  
 = electrons

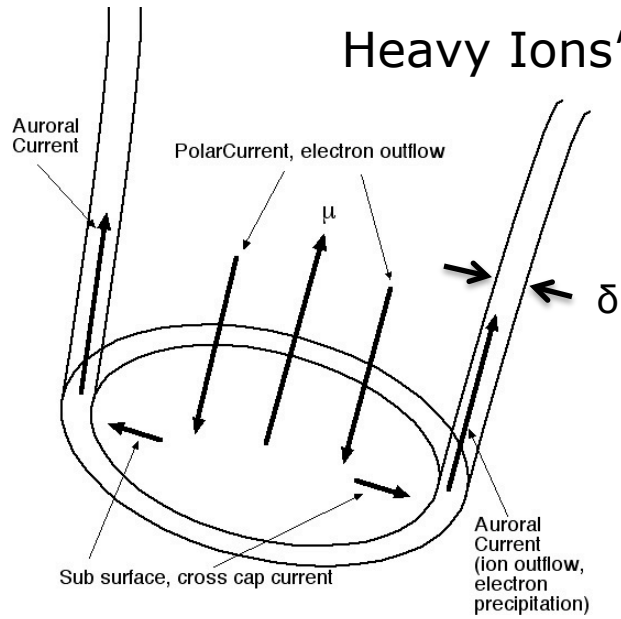


$E_{||}/B - B$  outside  
 current sheet,  
 $B=0$  inside

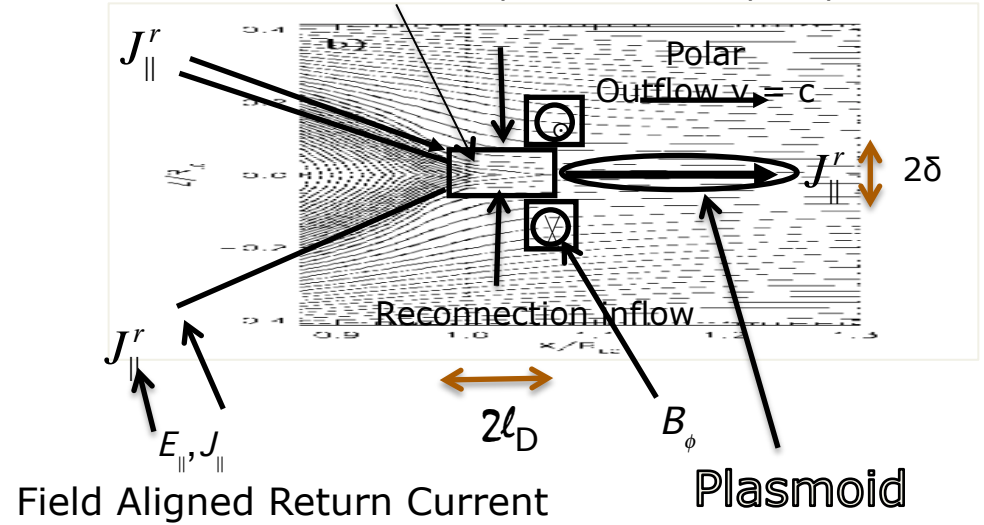
net ion outflow



# Heavy Ions' Source<sub>(14-17)</sub>



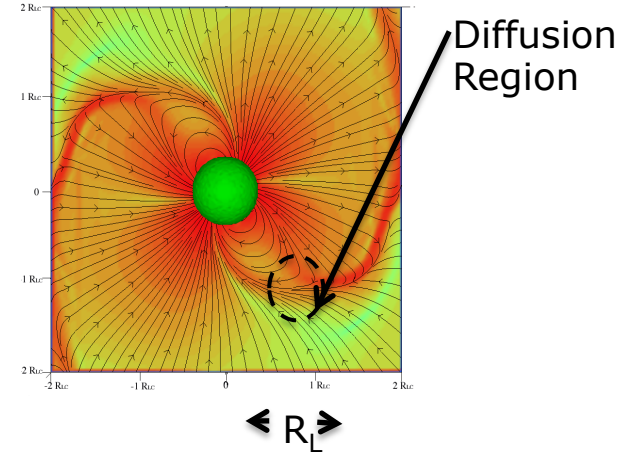
Unmagnetized Diffusion region & CS center: fed by reconnection inflow from wind thermal evaporation forms precipitation beam



Electron precipitation density  $n_* = I/r_{cap} \delta \gg GJ$   
 attracts upward ion beam from upper atmosphere ( $p^+, He^+, CNO^+?$ ) Or ocean ( $Fe^{+many?}$ )  
 Electron precipitation current  $\ll I = c\Phi \approx I_{ion}$

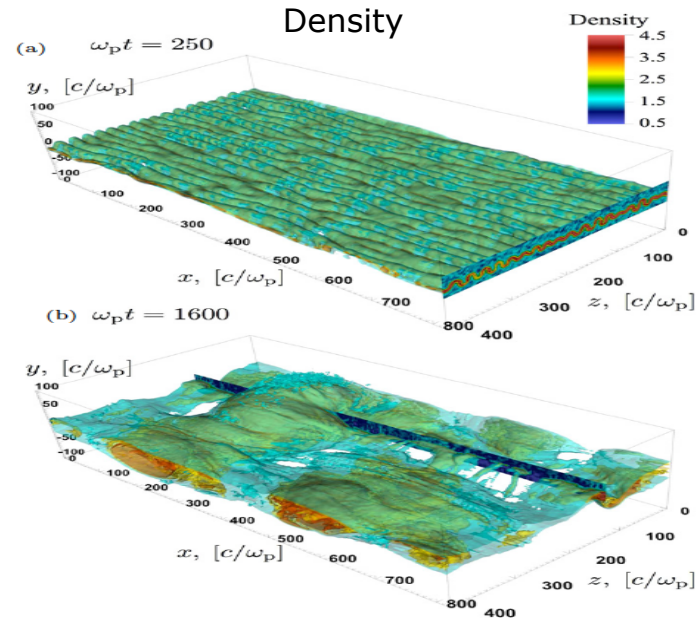
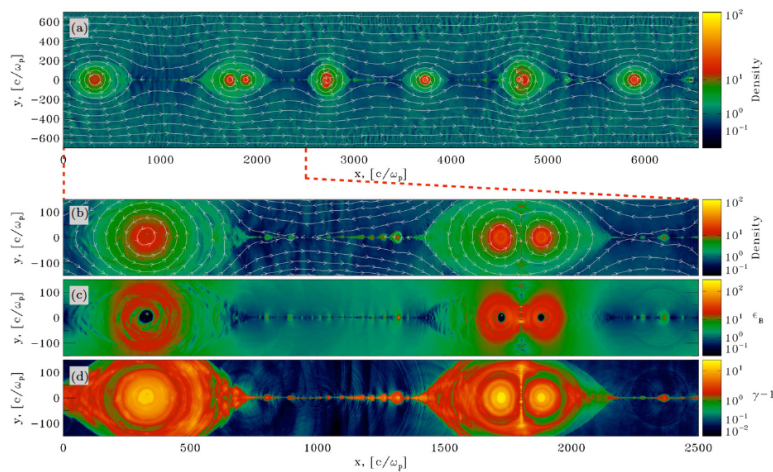
Acceleration along radial ( $r \gg R_L$ ) X-lines of reconnecting current sheet  $\sim$  linear accelerator

## Magnetic Y-line

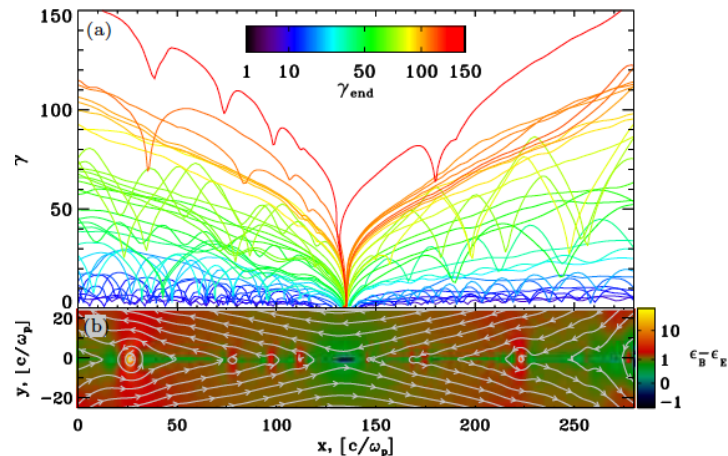
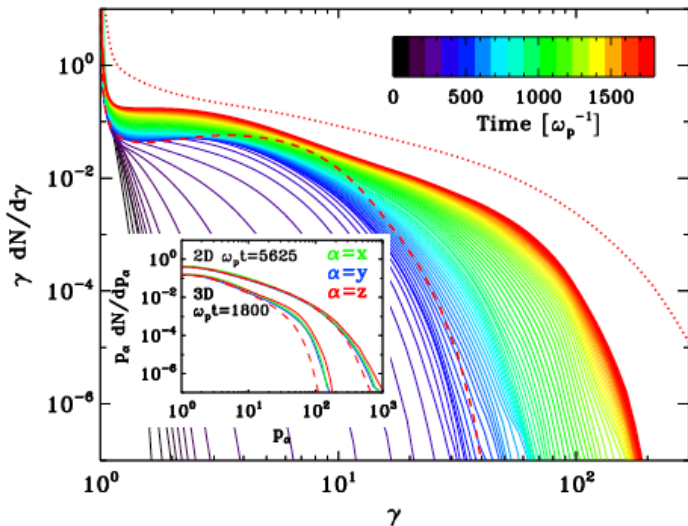


Electric return **current channel**  
 $\Omega \cdot \mu > 0$  Downward electron beam, upward ion beam  
 $\Omega \cdot \mu < 0$  Downward positron beam, upward electron beam

# 3D PIC e<sup>±</sup> (Sironi2014+)



Spectrum (whole box – radial, height, width  $\sim 6000 r_L \ll r_{wind} \sim 10^9 R_L$ )



Energy Histories

$dN/d\gamma \sim \gamma^{-1.3}$  –  $\sim$  monoenergetic, highest energy particles have most energy

Linear Accelerator = current sheet in wind; carries electric return current

Particle rate:

$$\dot{N}_i(t) = \frac{I_{return}(t)}{Ze} = \frac{\Omega^2(t)\mu}{Zec} \propto \frac{\Omega_i^2}{1 + \frac{t}{t_{EM}(\Omega_i)}}$$

Decays as star spins down (EM after initial 10s, neutrino heated wind gone):

$$t_{EM} = \frac{I_M c^3}{2\Omega^2 \mu^2} = \frac{10 \text{ yr}}{\mu_{30}^2} \left( \frac{P}{1 \text{ msec}} \right)^2$$

Maximum Energy: radial electric field  $E_r = \text{reconnection} = (v_{rec}/c)B_\phi$

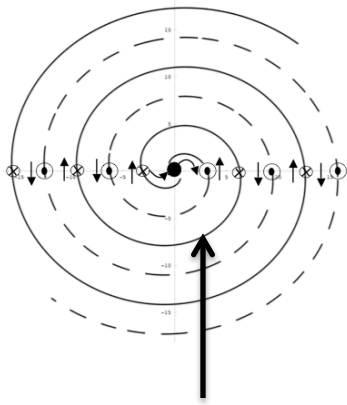
$$B_\phi = \Phi/r, \Phi = \mu\Omega^2/c^2 = 1.3 \times 10^{19} \mu_{30}/P_{\text{msec}}^2 \text{ Volts (magnetar: } \mu_{30} \sim 10^3)$$

$V_{rec}$  (simulations; simple 2 fluid theory) =  $0.8 \pm 0.2 v_A$ ,

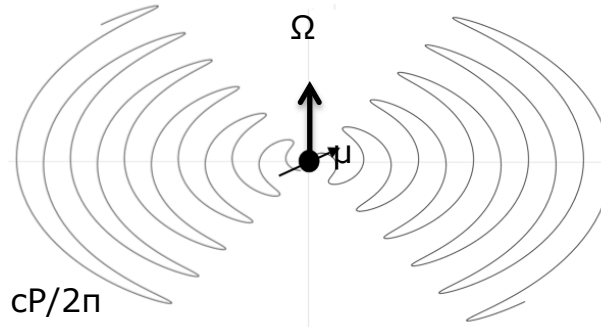
$v_A = c$  after initial 10s

$$\varepsilon = \gamma mc^2 = Ze \int_{r_{\min}}^{r_{\max}} E_r dr = Ze \left( \frac{v_{rec}}{c} \right) \int_{r_{\min}}^{r_{\max}} \frac{\Phi}{r} dr = Ze \left( \frac{v_{rec}}{c} \right) \Phi \ln \left( \frac{r_{\max}}{r_{\min}} \right)$$

# Oblique Rotators: Inner Wind Magnetically Striped



Equatorial cross-section



Meridional cross-section

$$\text{Wavelength} = R_L = cP/2\pi$$

Current Sheet

Dissipated in Wind Zone if

$$\Gamma_{\text{wind}} \leq \sigma_0 = \sqrt{\frac{\dot{E}_R}{Mc^2}} \ll 10^6$$

for Crab Nebula,  $\sigma_0 \sim 10^{3-4}$

Suggests stripes gone outside

$$r = R_{\text{diss}} = r_{\text{min}} \sim 10^{6-7} R_L = 10^{-(3-2)} R_{\text{TWS}}$$

$R_{\text{TWS}}$  = wind termination radius =  $r_{\text{max}}$

Mass loading of millisecond PSR = ?

Outside  $R_{\text{diss}}$ , current sheet flat = linear accelerator to  $R_{\text{TWS}}$

$$\varepsilon_{\text{max}}(t) = \gamma mc^2 = Ze \int_{r_{\text{min}}}^{r_{\text{max}}} E_r dr = Ze \left( \frac{v_{\text{rec}}}{c} \right) \int_{r_{\text{min}}}^{r_{\text{max}}} \frac{\Phi}{r} dr = Ze \left( \frac{v_{\text{rec}}}{c} \right) \Phi(t) \ln \left( \frac{r_{\text{max}}}{r_{\text{min}}} \right)$$



$R_{\text{diss}}$  = stripe dissipation radius

$10^7 R_L$  based on fast reconnection of striped current sheet

Site of Crab gamma ray flares? –

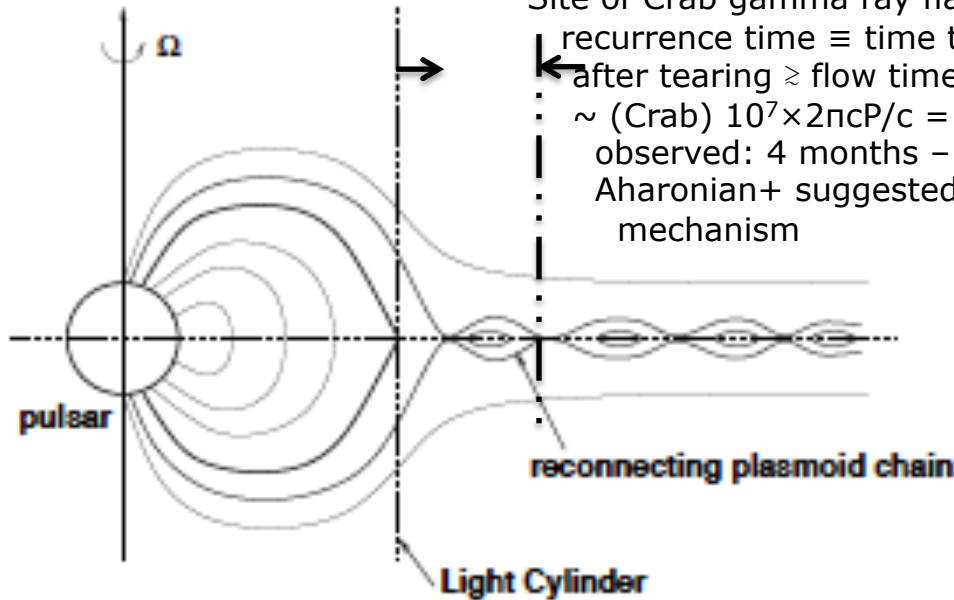
recurrence time  $\equiv$  time to restore current sheet

← after tearing  $\geq$  flow time from LC =  $R_{\text{diss}}/c$

$\sim$  (Crab)  $10^7 \times 2\pi c P/c = 0.9$  months –

observed: 4 months – 1 year

Aharonian+ suggested dissipation at  $30 R_L$ , without mechanism



$$R_{\text{TWS}} \sim 10^9 R_L \text{ (Crab)}$$

Particle Spectrum  $\propto E^{-1.3}$  from one star too hard for UHECR

Heavy ion source appealing, so

Superpose many stars/galaxies with a spectrum of voltages? (Kotera)

Process heavy ions (Fe?) in SNe ejecta shell? (Fang+)

# BLOWOUT

possible relation to hypernova models of GRBs -  
 fast rotating magnetic core forms in core collapse supernova;  
 magnetic pressure explodes stellar envelope, Compton  
 upscatter of radiation field as wind escapes creates (slow) GRB?

Rare compact objects:  $v_m = 10^{-4} v_{m4} \text{ yr}^{-1}$  Suggests unusual  
 core collapse SNe - Ib/c?

Newly formed magnetic core dumps  
 EM energy (B fields,...) in a few  
 minutes - initial spin down by GW  
 emission,

$$t_{GR}(\Omega_i) \approx \frac{30}{\Omega_4^4 (\epsilon / 10^{-2})^2} \text{ sec},$$

$$\Delta E_{EM}(t < t_{GR}) \approx 5 \times 10^{51} \left( \frac{\mu_{33}}{\epsilon / 10^{-2}} \right)^2 \text{ ergs} = 0.1 \left( \frac{1}{2} I \Omega_i^2 \right) \quad (\Omega_i \sim 10^4 \text{ s}^{-1})$$

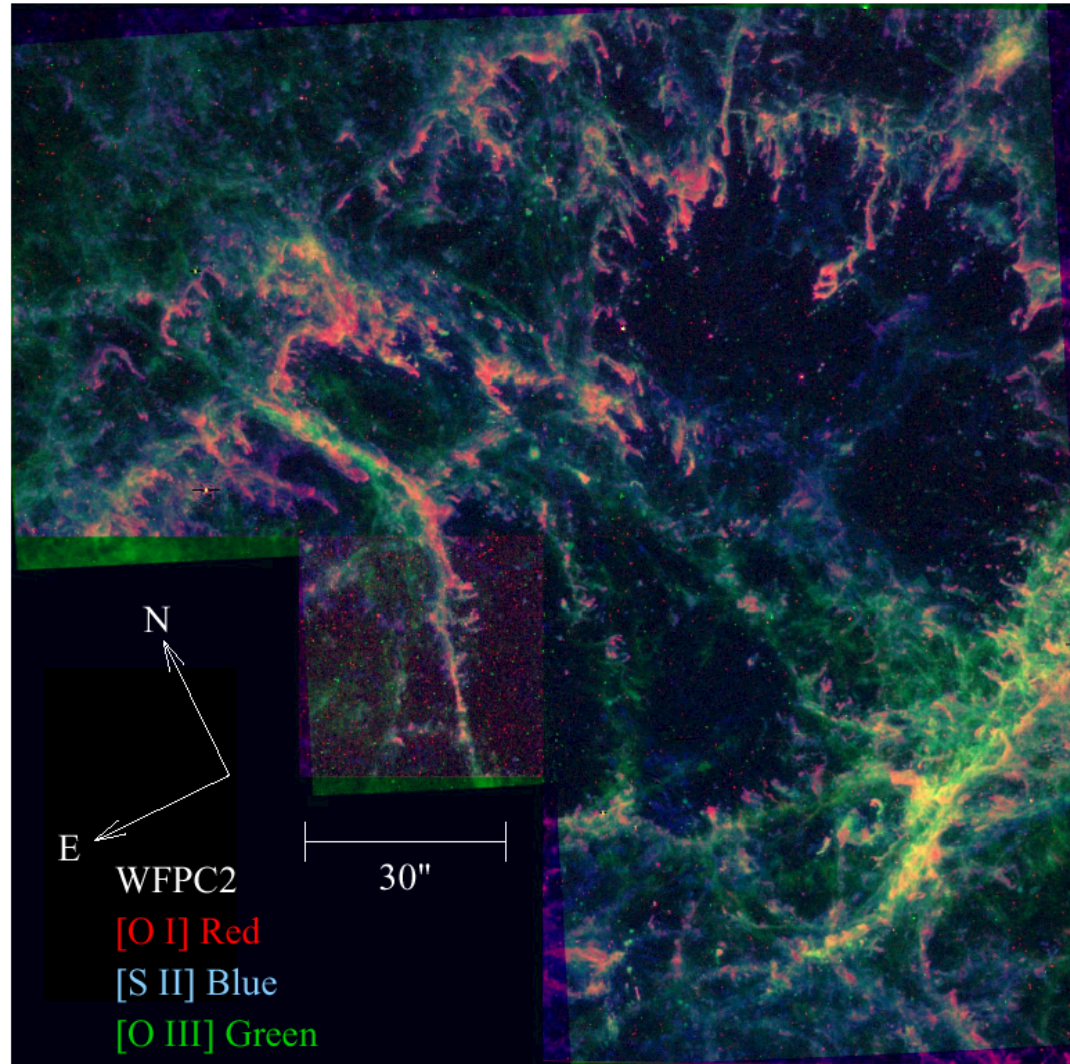
Pre SN star likely compact, with fairly short dynamical time  
 Wheeler et al 2000 model:

$$t_d \sim 20 \frac{(R_* / 10^{5.5} \text{ km})^{3/2}}{(M / M_{sun})^{1/2}} \text{ sec}$$

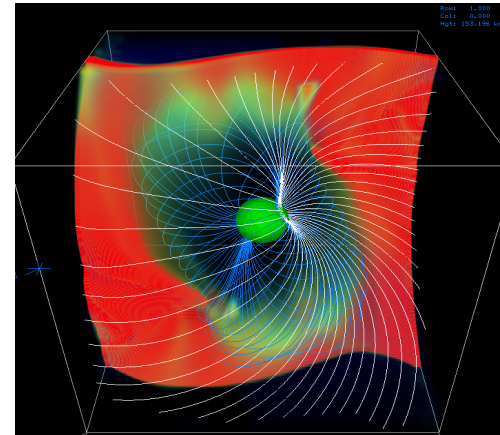
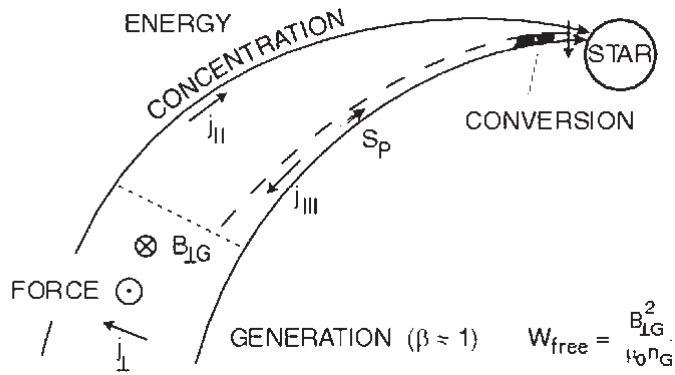
# Injected EM Energy disrupts pre SN envelope in the dynamical time

Rayleigh-Taylor of light EM energy shreds envelope in time  $t_{\text{dyn}}$ , short compared to standard SN,

Wind then expands freely, blows bubble in ISM, expansion non-relativistic after 1 year – deposits  $\sim 2-5 \times 10^{50-51}$  ergs/neutron \* in the ISM, limited by gravitational wave loss



Crab filaments - RT shredded ejecta (Sankrit Hester et al); also Gamma ray leakage from '87a



# THE END